

SOLAR INDICES BULLETIN

JUNE 1995

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◆ SOLAR RADIO EMISSIONS

The quiet sun emits radio energy with a slowly varying intensity. These radio fluxes, which stem from atmospheric layers high in the chromosphere and low in the corona, change gradually from day-to-day, in response to the number and size of spot groups on the solar disk. The table below gives daily measurements of this slowly varying emission at selected wavelengths between about 1 and 100 centimeters. Many observatories record quiet-sun radio fluxes at the same local time each day and correct them to within a few percent for factors such as antenna gain, bursts in progress, atmospheric absorption, and sky background temperature. At 2800 megahertz (10.7 centimeters) flux observations summed over the sun's disk have been made continuously since February 1947.

During low periods of solar activity, the flux never falls to zero, because the sun emits at all wavelengths with or without the presence of spots. The lowest daily Ottawa flux since 1947 occurred on November 3, 1954. On that day the observed noon value dropped to 62.6 units; the highest observed value of 457.0 occurred on April 7, 1947.

The preliminary observed and adjusted Penticton fluxes tabulated here are the "Series C" values reported by Canada's Dominion Radio Astrophysical Observatory in Penticton, British Columbia. Observed numbers are less refined, since they contain fluctuations as large as ±7% from the continuously changing sun-earth distance. Adjusted fluxes have this variation removed; they show the energy received at the mean distance between the sun and earth. Gaps in the Learmonth, Australia (LEAR) data reflect equipment problems. Fluxes measured either at Palehua on the Hawaiian Islands, or at San Vito, Italy, will be substituted for frequencies at which many Learmonth values are missing.

◆ SOLAR FLUX TABLE

Numbers in parentheses in the column headings below denote frequencies in megahertz. Each entry is given in solar flux units--a measure of energy received per unit time, per unit area, per unit frequency interval. One solar flux unit equals 10^{-22} J/m²Hzsec.

JUNE 1995 PRELIMINARY SUNSPOT NUMBERS AND SOLAR RADIO FLUX

MAY 1995 FINAL FLUX

Day	Sunspot Number	Obs Flux		Solar Flux Adjusted to 1 Astronomical Unit							
		Pentic (2800)	LEAR (15400)	LEAR (8800)	LEAR (4995)	Pentic (2800)	LEAR (2695)	LEAR (1415)	LEAR (610)	LEAR (410)	LEAR (245)
01	8	71	523	225	122	73	67	50	30	23	11
02	9	73	489	225	123	75	69	51	33	24	12
03	12	75	523	228	127	77	73	55	34	24	12
04	14	79	514	227	128	81	73	56	34	24	11
05	30	82	511	235	135	84	81	60	37	26	14
06	24	79	518	230	134	82	81	61	37	28	15
07	28	83	518	231	133	85	82	63	35	26	18
08	28	84	529	229	133	87	81	64	40	28	21
09	24	86	518	231	135	89	84	65	40	27	12
10	30	84	504	217	135	87	84	66	40	28	13
11	24	83	513	232	133	85	82	64	39	28	12
12	20	81	517	228	131	83	81	61	38	26	13
13	17	77	485	224	130	79	80	59	35	25	13
14	12	76	520	226	128	78	76	57	35	25	13
15	0	73	516	228	127	76	75	56	34	25	14
16	0	71	513	226	125	73	71	53	34	24	12
17	7	70	534	223	123	72	69	51	33	24	12
18	9	70	503	224	123	72	70	51	32	23	11
19	18	71	497	221	123	74	70	52	33	24	12
20	13	72	507	223	124	74	72	54	31	22	12
21	15	72	377	204	117	75	69	52	15	10	5
22	17	73	509	221	122	76	72	52	32	22	11
23	16	72	430	219	121	74	71	52	31	22	12
24	13	71	—	—	—	73	—	—	—	—	—
25	13	71	519	218	121	73	71	52	32	23	13
26	12	71	518	221	122	73	69	52	33	23	11
27	10	72	518	221	122	75	71	53	34	24	12
28	9	75	523	222	124	77	71	53	33	24	11
29	14	78	539	225	126	80	74	54	34	24	12
30	27	78	548	232	131	81	77	57	35	25	12
Mean	16	76	508	225	127	78	75	56	34	24	12

Observed Pentic (2800)	Adjusted Pentic (2800)
68.6	69.7
68.8	69.9
69.0	70.2
70.6	71.8
72.9	74.2
76.4	77.8
77.6	79.0
78.0	79.4
77.8	79.2
77.8	79.3
76.8	78.4
80.8	82.4
80.5	82.2
79.9	81.6
85.9	87.8
93.8	95.9
95.2	97.3
91.3	93.4
86.3	88.3
80.8	82.8
75.1	77.0
70.8	72.6
68.7	70.4
67.1	68.8
67.1	68.9
66.3	68.1
65.8	67.5
66.6	68.4
66.8	68.7
67.5	69.4
68.9	70.9
75.5	77.1

◆ SUNSPOT COUNTS

In 1848 the Swiss astronomer Johann Rudolph Wolf introduced a daily measurement of sunspot number. His method, which is still used today, counts the total number of spots visible on the face of the sun and the number of groups into which they cluster, because neither quantity alone satisfactorily measures the level of sunspot activity.

An observer computes a daily sunspot number by multiplying his estimated number of groups by ten and then adding this product to his total count of individual spots. Results, however, vary greatly, since the measurement strongly depends on observer interpretation and experience and on the stability of the earth's atmosphere above the observing site. Moreover, the use of earth as a platform from which to record these numbers contributes to their variability, too, because the sun rotates and the evolving spot groups are distributed unevenly across solar longitudes. To compensate for these limitations, each daily international number is computed as a weighted average of measurements made from a network of

cooperating observatories. The international sunspot numbers tabulated on page 1 are provisional values taken from a bulletin prepared monthly by Pierre Cugnon of the SUNSPOT INDEX DATA CENTER, 3 avenue Circulaire, B-1180 BRUXELLES, BELGIUM. The June 1995 data combine observations from 44 stations.

◆ HISTORICAL SUNSPOT COUNTS

How do sunspot numbers in the table on page 1 compare to the largest values ever recorded? The highest daily count on record occurred December 24-25, 1957. On each of those days the sunspot number totaled 355. In contrast, during years near the spot cycle minimum, the count can fall to zero. Today, much more sophisticated measurements of solar activity are made routinely, but none has the link with the past that sunspot numbers have. Our archives, for example, include reconstructed daily values from January 8, 1818; monthly means from January 1749; and yearly means beginning in 1700.

SMOOTHED (OBSERVED AND PREDICTED) SUNSPOT NUMBERS: CYCLES 21, 22 AND 23

Year	Jan	Feb	Mar	Apr	May	Jun	Jul	Aug	Sep	Oct	Nov	Dec	Mean
1985	20	20	19	18	18	18	17	17	17	17	17	15	18
1986	14	13	13	14	14	14	14	13	12*	13	15	16	14
1987	18	20	22	24	27	28	31	35	39	44	47	51	32
1988	58	65	71	78	84	94	104	114	121	125	130	138	98
1989	142	145	150	154	157	158	159	158	157	157	158	154	154
1990	151	153	152	149	147	144	141	141	142	142	142	144	146
1991	148	148	147	147	146	145	146	147	145	142	138	132	144
1992	124	115	108	103	100	97	91	84	80	76	74	73	94
1993	71	69	67	64	60	56	55	52	48	45	41	38	56
1994	37	35	34	34	33	31	29	27	27	27	26	26	30
1995	25 (4)	24 (5)	23 (7)	22 (8)	20 (8)	20 (9)	19 (10)	18 (11)	18 (11)	17 (11)	17 (11)	16 (10)	20 (9)
1996	15 (10)	13 (11)	13 (11)	12 (12)	12 (12)	11 (12)	11 (12)	10 (12)	10 (12)	9 (12)	9 (12)	8 (11)	11 (12)
1997	9 (11)	9 (11)	9 (10)	9 (10)	10 (11)	11 (12)	11 (13)	12 (15)	13 (16)	13 (17)	14 (19)	15 (22)	11 (14)

Sep 1986 marks Cycle 21's minimum and the onset of cycle 22, which reached a maximum in July 1989.

◆ SUNSPOT NUMBER PREDICTIONS

For the end of Solar Cycle 21, and the beginning of Cycle 22, the table gives smoothed sunspot numbers up to the one calculated that first uses the most recently measured monthly mean. These smoothed, observed values are based on final, unsmoothed monthly means through March 1995 and on provisional ones thereafter. We compute a smoothed monthly mean by forming the arithmetic average of two sequential 12-month running means of monthly means.

Table entries with numbers in parentheses below them denote predictions by the McNish-Lincoln method. This method estimates future numbers by adding a correction to the mean of all cycles that is proportional to the departure of earlier values of the current cycle from the mean cycle. (See page 9 in the July 1987 supplement to *Solar-Geophysical Data*). We use and predict only smoothed monthly means, because we believe the errors are too great to

estimate any values more precise. In the table above, adding the number in parentheses to the predicted value generates the upper limit of the 90% confidence interval; subtracting the number from the predicted value generates the lower limit. Consider, for example the December 1995 prediction. There exists a 90% chance that in December 1995 the actual smoothed sunspot number will fall somewhere between 6 and 26.

The McNish-Lincoln prediction method generates useful estimates of smoothed, monthly mean sunspot numbers for no more than 12 months ahead. Beyond a year these predictions regress rapidly toward the mean of all 13 cycles used in the computation. Moreover, the method is very sensitive to the date defined as the beginning of the current sunspot cycle, that is, to the date of the most recent sunspot minimum. The new cycle predictions tabulated above are based on the minimum value of 12.3 that occurred in September 1986.

Although every effort has been made to ensure that these data are correct, we can assume no liability for any damages their inaccuracies might cause. The charge for a 1-year subscription to this monthly bulletin in US\$21.00. To become a subscriber, you may either call (303) 497-6346 or write the NATIONAL GEOPHYSICAL DATA CENTER, Solar-Terrestrial Physics Division (E/GC2), 325 Broadway, Boulder, Colorado 80303 USA. Please include with your written order a cheque or money order payable in U.S. currency to the "Department of Commerce, NOAA/NGDC". Payment may also be made through VISA, MasterCard or American Express credit cards.